

The dynamics behind Titan's methane clouds

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We present results of an axisymmetric global circulation model of Titan with a simplified suite of atmospheric physics forced by seasonally varying insolation. The recent discovery of midlatitude tropospheric clouds on Titan has caused much excitement about the roles of surface sources of methane and the global circulation in forming clouds. Although localized surface sources, such as methane geysers or “cryovolcanoes,” have been invoked to explain these clouds, we find in this work that clouds appear in regions of convergence by the mean meridional circulation and over the poles during solstices, where the solar forcing reaches its seasonal maximum. Other regions are inhibited from forming clouds because of dynamical transports of methane and strong subsidence. We find that for a variety of moist regimes, i.e., with the effect of methane thermodynamics included, the observed cloud features can be explained by the large-scale dynamics of the atmosphere. Clouds at the solstitial pole are found to be a robust feature of Titan's dynamics, whereas isolated midlatitude clouds are present exclusively in a variety of moist dynamical regimes. In all cases, even without including methane thermodynamics, our model ceases to produce polar clouds ≈ 4 –6 terrestrial years after solstices.

atmospheres | climate | planetary science

Methane clouds were first observed in Titan's troposphere as occasional brightening in disk-averaged infrared photometry (1). Clouds were observed in the southern hemisphere later during southern spring (2). In 2001, clouds were observed near the South Pole, presumably caused by strong solar heating at the pole just preceding Southern summer solstice (SSS, which occurred in October 2002) (3, 4). The first observations of midlatitude clouds after SSS were reported in 2005 (5) and confirmed by Cassini observations (6, 7) at a time when Titan was progressing toward Southern autumnal equinox. Efforts have been focused on explaining what controls the position of clouds and mechanisms that set the timing of observed shifts in their positions. Earth-based observations show midlatitude clouds preferentially form at a particular latitude and longitude, which suggests they owe their existence to localized surface sources of methane (8). But, the longitudinal distribution of clouds is more uniform if both ground-based and Cassini data are taken into account (7). It is known from ground-based observations that large outbursts in cloud activity occasionally occur (9). Ground-based observations from late 2005 revealed that Titan's south polar clouds, the most persistent cloud features, have completely dissipated, and concurrently the sporadic activity at midlatitudes first subsided and then re-established at lower latitudes (10).[†]

The position, variability, and seasonality of clouds is certainly, in some manner, influenced by the seasonally varying distribution of solar heating. However, a local radiative-convective model with no large-scale dynamics will produce convection at essentially all latitudes (see Fig. 3, which is published as supporting information on the PNAS web site). On Earth the pattern of tropical convection is strongly modulated by large-scale dynamics. In particular, moisture converged by the large-scale dynamics condenses out within the upward branch of the mean meridional circulation (MMC), releasing latent heat there.

This heating tends to narrow and strengthen the updraft and produce persistent clouds and precipitation in a region known as the Intertropical Convergence Zone (ITCZ). The upward flow in the ITCZ is balanced by large areas of subsiding dry air that suppress convection and diverge moisture at the surface, which produce the subtropical deserts (11). There is every reason to believe that dynamics have an equally profound effect on convection and cloud formation on Titan because methane plays an equivalent role to water vapor in Earth's climate. Because of its slow rotation rate (1/16 of Earth's rotation rate), theory predicts Titan's MMC should extend to the poles, and temperatures should be roughly uniform across the region of overturning (12); these features are borne out by the weak equator-to-pole surface temperature gradient observed by Voyager (13). Titan, in effect, is an “all tropics” planet, and we are justified in interpreting, at least some level, cloud patterns in the framework of existing theories of tropical dynamics.

Titan's season is ≈ 30 Earth years long; because the SSS was in 2002, the Southern autumnal equinox will not happen until late 2009. Until recently, the latitudinal positions of midlatitude clouds seemed not to change significantly with seasonal shifts in maximum solar heating, at least for the last few years during which disk-resolved observations are available (8). Modeling studies of the response of Titan's troposphere to seasonally varying forcing could shed light on the mechanisms controlling seasonal cloud patterns. General circulation models (GCMs) of Titan to date have focused mainly on stratospheric dynamics. The stratosphere of Titan is superrotating (global angular momentum exceeds solid-body rotation) and contains a thick organic “haze,” which shrouds the lower atmosphere.

Stratospheric zonal wind velocities over the equator are 100 m/s, and the midlatitude jets reach ≈ 150 m/s (14). GCMs with Titan's rotation rate have shown that eddies converge momentum toward the equator to sustain superrotating winds (15–18); this mechanism was first proposed to explain superrotating winds on Venus (19). Simulations with coupled interaction between the dynamics of the stratosphere and the organic haze production and distribution can reproduce the seasonality of latitudinal contrasts in haze thickness, which explains the polar “hooding” seen in observations (6, 20–22).

The discovery of tropospheric methane clouds in the midlatitudes of the summer hemisphere makes a strong case for a detailed study of the response of tropospheric dynamics to seasonal forcing. Indeed, interest in modeling Titan's troposphere is on the rise. A recent GCM simulation was able to

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Abbreviations: MMC, mean meridional circulation; SSS, Southern summer solstice; GCM, general circulation model; ITCZ, Intertropical Convergence Zone; LSC, large-scale condensation.

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reproduce the basic features of current cloud patterns; in that model, clouds are produced through methane transports by the MMC and symmetric eddies. This study predicted clouds at both poles at all seasons and midlatitude clouds in both hemispheres around both solstices and in general showed extremely weak seasonality (23). Another modeling study of Titan's troposphere focused on the latent feedback of methane condensation during different seasons, but did not account for the full seasonal cycle (24). The meteorological impact of surface types, primarily surface thermal inertia, has also been explored in a GCM; this study primarily addressed seasonal changes in surface temperatures, but also concluded the distribution of observed cloud features is consistent with a solid surface with low thermal inertia (25). A GCM simulation was also used to predict balloon trajectories for a proposed mission to Titan (26). The basic behavior of the tropospheric system, though, has hardly been explored, and in particular the relevant time scales in the system require further study.

The radiative damping time of Titan's troposphere is very long, similar in magnitude to the Titan season. It might be expected, therefore, that seasonality in Titan's tropospheric circulation and precipitation would be significantly attenuated. However, because there is no global ocean on Titan (27) the surface presumably has quite low thermal inertia. Furthermore because the tropics (which span the globe of Titan) are thermally uniform, small perturbations in surface temperature from changes in the surface energy balance can have very dramatic effects on the position of convection (for a description of this effect in the Earth's tropics, see ref. 28). Because Titan's surface has very low heat capacity, we expect it to respond quickly to changes in solar forcing if the atmosphere is unable to communicate its heat capacity to the surface; the warmest surface temperatures will be found at the subsolar point in this case, and the pattern of convection/precipitation will follow along in-phase. Or if the atmosphere is able to communicate its large heat capacity to the surface by some mechanism, we expect the seasonal cycle of convection/precipitation to be somewhat attenuated.

Radiative and turbulent (sensible and latent) energy fluxes at the surface couple the surface and the atmosphere. Here we provide estimates of these energy fluxes for Titan's surface. The solar forcing at Titan is only a few W/m^2 , and the atmosphere and surface at $\approx 100 \text{ K}$ only emit $5\text{--}6 \text{ W/m}^2$ of thermal radiation. The available flux of latent energy at the surface can be estimated by an assumed turbulent mass flux^{||} $\times$ the mixing ratio of the condensable $\times$ its latent heat of vaporization. By this estimate, the available latent energy flux at Titan's surface is 10^2 to 10^3 W/m^2 . By the bulk formula, sensible heat flux is generally quite weak compared with the latent heat flux. In summary, latent fluxes have the potential to greatly exceed other energy fluxes, although they will be restricted by energy balance. We might expect, therefore, that the seasonal cycle of convection/precipitation will be sensitive to changes in surface latent energy flux because this energy has the greatest potential to couple the low-thermal-inertia surface to the high-thermal-inertia atmosphere.

Our goal is to provide a simple modeling framework that will allow us to constrain the physical mechanisms affecting the seasonal behavior of Titan's clouds. Observations of the surface of Titan have shown vast areas of typical desert morphologies such as dunes and other areas that may be densely populated with lakes of methane (see public images available at <http://saturn.jpl.nasa.gov>). As noted previously, there is no evidence

for a global ocean. In the present study we explore the effect various levels of methane concentrations in the troposphere, as modified by methane convection and the supply of methane from the surface, have on the dynamics of the troposphere. We find that a simple gray radiation scheme is able to reproduce the important features of the observed temperature profile (see *Methods*). Such a computationally efficient scheme allows us to thoroughly explore parameter dependencies of our model and sets the stage for us to extend our parameter study to three dimensions, which would otherwise be prohibitive because of the computational time required for the full radiation scheme. In the next section, we establish consistency of our model with observations. We then present our model results and offer some conclusions. For an overview of our model see *Methods*.

Consistency with Observations

Given the recent Cassini/Huygens observations of Titan's troposphere, it is important to establish consistency of our model simulations with the available data. Our dynamical simulations reproduce the observed weak equator-to-pole surface temperature gradient of a few Kelvin. Radiative-convective simulations, in contrast, show a much larger surface temperature gradient of between 10 and 20 K, so the MMC is strongly reorganizing the thermal structure in the troposphere, as expected from theory (see Introduction) (12). The diagnosed tropospheric vertical temperature profiles in simulations with the effect of methane thermodynamics included are consistent with observations by the Huygens Atmospheric Structure Instrument, whereas the lapse rate in the dry simulation is generally somewhat higher than is observed (29).

Observations of the shearing of Titan's methane clouds reveal typical zonal wind speeds of $\mathcal{O}(10 \text{ ms}^{-1})$ (6), which are similar to zonal winds produced in our simulations. The Doppler Wind Experiment on the Huygens probe, which descended at $\approx 10^\circ\text{S}$ latitude, measured easterlies in the boundary layer and weak westerlies aloft (30). The MMC in our model is able to sustain angular-momentum-conserving zonal winds aloft because the dynamical overturning of our model is short compared with a seasonal cycle and because we have no parameterization for horizontal momentum mixing. There is some question as to whether eddies mix momentum up-gradient and produce superrotation in the troposphere, as they do in the stratosphere.** We do not attempt to model superrotation in our current model because tropospheric zonal wind measurements are too sparse to rule out the characteristic uniform angular momentum distribution of a Held-Hou-like MMC (12).

Our moist^{††} convection scheme relaxes methane to a specified relative humidity when convectively unstable. We therefore do not expect our diagnosed methane abundances to match observations in general. This potential discrepancy is tolerable for our purposes because we do not seek to account for the radiative feedback of methane, and it has been shown that Earth's dynamics are relatively insensitive to the choice of this relative humidity parameter, although it does have an effect on the surface energy budget (31). Our intermediate-moisture, dynamical simulation best captures the essence of the observed vertical profile of methane abundances at low latitudes during the season when the Huygens probe gathered data (32).

Results: Model Sensitivity to Methane Vapor Concentration

To study the feedback of methane thermodynamics on Titan's tropospheric dynamics, we performed three simulations with

^{||}We estimate the mass flux by using the bulk aerodynamic formula $C_d \rho U_0$, where C_d is a roughness factor, ρ is the density of air in the lower atmosphere, and U_0 is a gustiness parameter.

**There is some evidence that zonal winds reach 30 ms^{-1} at midlatitudes, which is greater than a momentum-conserving gradient wind allows.

^{††}Here and throughout, we refer to methane vapor (or the methane cycle) in the troposphere as moisture, in analogy to the cycle of water vapor in the Earth's troposphere.

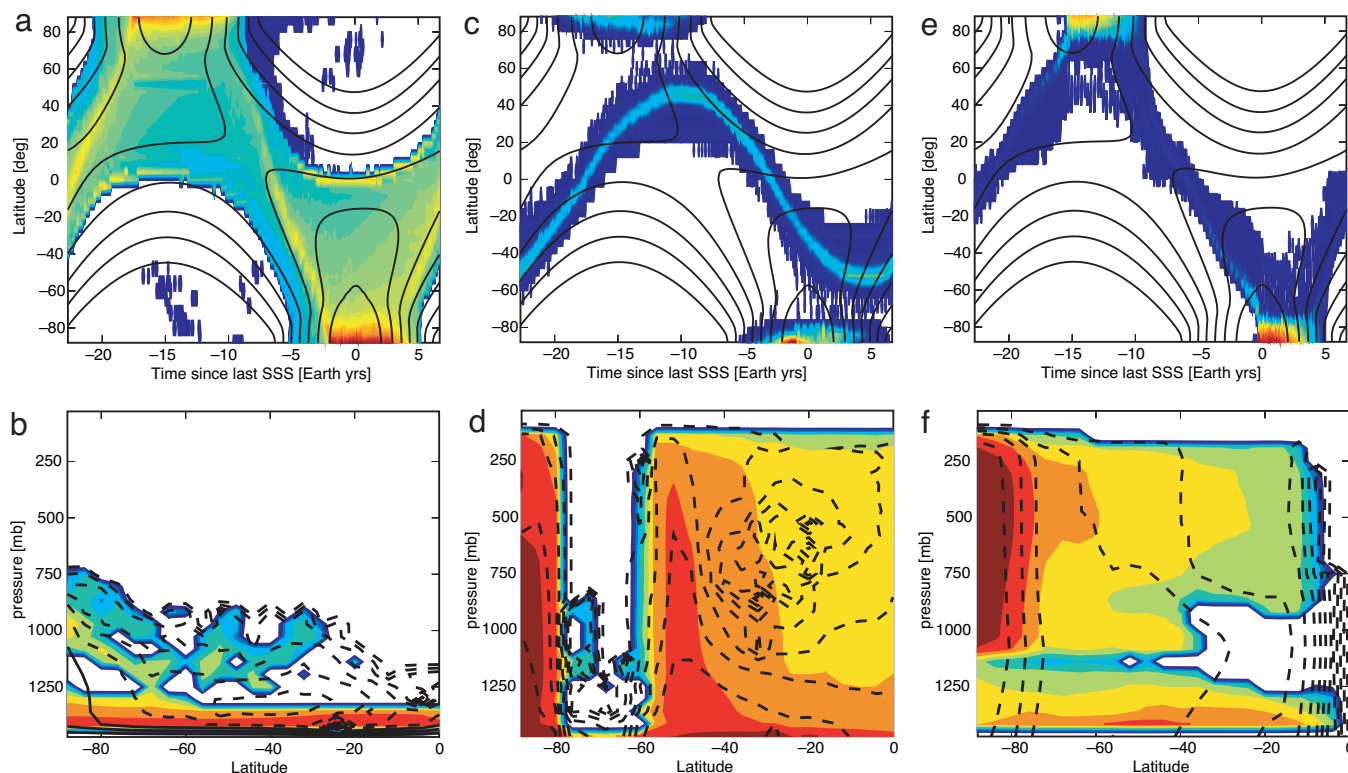


Fig. 1. Model precipitation and convection. (a, c, and e) Contour plots of precipitation (filled contours) for our three simulations, with the pattern of solar forcing at the surface overlaid (black lines) for reference. (a) Dry seasonality. (c) Moist seasonality. (e) Intermediate seasonality. (b, d, and f) Contour plots of the logarithm of the averages of convective perturbations for the 10 terrestrial years bracketing SSS; filled contours are the convective heating rate in K per day on the same color scale (10^{-6} , cool colors to $10^{-1.5}$, warm colors) and dashed contours are the convective drying rate in g/kg per day. “Stepping” patterns at the edges of contours are representative of the resolution of our simulations. (b) Dry convection. (d) Moist convection. (f) Intermediate convection.

varying amounts of methane vapor (see *Methods* for details). Fig. 1 a, c, and e shows contour plots of precipitation for our three dynamical simulations over one Titan season with solar forcing at the surface overlaid for reference; we infer clouds in our simulations by the presence of condensation, which always produces precipitation in our convection scheme (snapshots of the states of our model atmospheres can be found in Fig. 4, which is published as supporting information on the PNAS web site). In Fig. 1 b, d, and f we show averages of southern hemisphere convective temperature and methane concentration tendencies (in K per day and g/kg per day, respectively) for the 10 terrestrial years bracketing SSS, to show where convection has the potential for creating clouds. Below we analyze each case in turn within the theoretical framework outlined in the Introduction.

In the dry case, we remove the latent heating effects of methane evaporation/condensation (see *Methods*). In this limit, methane is advected by the dynamics and condenses in the atmosphere to produce precipitation without affecting the thermal structure. The dry case should not be thought of as representing an actual planet, but by comparison to moist simulations serves to illustrate the effect of methane thermodynamics on the circulation of the atmosphere. The most obvious feature in the dry case is that of widespread precipitation in the summer hemisphere, as seen in Fig. 1a. Also notice precipitation is in-phase with the solar forcing, so the thermal inertia of the atmosphere is having little damping effect on the seasonal cycle. During solstices, the ITCZ reaches the summer pole and the MMC extends to midlatitudes in the winter hemisphere, and it switches directions with seasons. Fig. 1b shows the areas of convection are generally shallower than in the moist cases, which is also to be expected in the absence of latent heat effects by

methane and is seen in radiative-convective equilibrium (see Fig. 3); the heights associated with the pressure levels at the top of the convection are shallower than the observed convective clouds at midlatitudes (7). Given the widespread precipitation predicted in the summer hemisphere, the dry case is inconsistent with observations.

The moist case shows a drastically altered precipitation pattern. Dynamical interactions with latent heating are clearly having a profound effect on convective cloud formation, suppressing convection throughout much of the atmosphere, as can be seen by comparing the dry and moist runs in Fig. 1 a and c. First, notice the meandering precipitation band that extends from midlatitudes in one hemisphere to the other and is out-of-phase with the solar forcing. This feature is the ITCZ and marks the point of surface convergence of the MMC and resulting large-scale updraft. The lag in phase is caused by the strong thermal coupling of the surface and atmosphere caused by surface fluxes, which allows the atmosphere to communicate its large thermal inertia to the surface and dampen changes in surface temperatures. The ITCZ reaches its maximum poleward extent ≈ 5 terrestrial years after solstices. Methane also significantly destabilizes the lower atmosphere toward convection; Fig. 1d shows the convection reaches much deeper into the atmosphere in the moist cases vs. the dry cases. The ITCZ appears in Fig. 1d as the vertical column of convective tendencies centered at 50° S latitude, and the extended tendencies northward of this column; the smearing in latitude is caused by the ITCZ just having reached its southernmost latitude by the end of the displayed 10-year average. The second precipitation feature is caused by a complex of polar convection that is symmetrically distributed in time around solstices and only at the summer pole,

Table 1. Summary of parameter settings and results for our model runs

Model	Convection humidity parameter, %*	Boundary layer humidity parameter, %†	Latent effects included‡	Consistent with observations§
Radiative-convective, dry	80	100	No	No
Radiative-convective, moist	80	100	Yes	No
Dynamical, dry	80	100	No	No
Dynamical, moist	80	100	Yes	Yes
Dynamical, intermediate	40	50	Yes	Yes
Dynamical, moist LSC only	100	100	Yes	Maybe

Radiative-convective simulations are presented in *Supporting Text*, which is published as supporting information on the PNAS web site.

*The value of relative humidity the convection scheme requires for convection to occur.

†The value of relative humidity the boundary layer scheme attempts to relax toward.

‡Indicated is whether the latent effects of methane condensation and evaporation were included, which represents the primary distinction between dry and moist or intermediate models.

§Our judgment as to whether the results were consistent with available cloud observations.

as seen in Fig. 1c. Fig. 1d shows the most persistent strong convection during SSS is directly at the South Pole. Such clouds dissipate ≈ 5 years after solstices, about the time the ITCZ halts its march toward the pole. The moist case is found to be consistent with observations.

The intermediate case, in which the boundary layer scheme relaxes the first model layer toward 50% relative humidity and convection can occur in the interior at 40% relative humidity (see *Methods*), shows some features of both the dry and moist cases. As in the moist case, Fig. 1e shows a sparse precipitation pattern; dynamics are having a comparably profound effect on the distribution of clouds. But the ITCZ shows stronger seasonality like the dry case; the ITCZ is now able to reach the pole, and its phase lag with solar forcing is shortened to 2 terrestrial years. The lack of distinct columns of convective perturbations in Fig. 1f shows the ITCZ and polar convective patterns have become the same feature. The migration of the ITCZ smears out the convective patterns seen in the 10-year average. It is only when the ITCZ reaches the poles that persistent precipitation appears at the summer pole. During this time, variability in the circulation triggers sporadic cloud activity at mid and high latitudes, as can be seen by the streaks of precipitation that move equatorward in Fig. 1e. At ≈ 5 years after solstices, polar clouds cease and the ITCZ begins to migrate back toward the equator. The intermediate case is also found to be consistent with observations of clouds. In particular, the migration pattern of precipitation during and after SSS is quite similar to the distribution of observed clouds during this same time (10).

Conclusions

We have performed simplified, axisymmetric simulations of the dynamics of Titan's troposphere with realistic seasonal forcing and methane thermodynamics. Polar clouds are found to be a robust feature of Titan's solstices. In all cases presented, polar clouds cease ≈ 4 –6 terrestrial years after solstices. Isolated midlatitude clouds near solstices are produced only in simulations with methane thermodynamics included, because in these cases the ITCZ is enhanced and narrowed, methane is converged to the updraft of the MMC at the surface, and broad regions of large-scale subsidence suppress convection over much of the globe. Our model produces a seasonally varying ITCZ that is in-phase with solar forcing in the dry case and out-of-phase with solar forcing when methane thermodynamics is included. We find the magnitude of the seasonal response, measured both as the phase lag and the maximum latitude reached by the ITCZ, is primarily constrained by the amount of methane being supplied from the surface. Latent heat fluxes from evaporation of methane at the surface stiffens the system against seasonal

forcing, because the atmosphere is able to effectively communicate its large thermal inertia to the surface. In the absence of latent fluxes, the surface warms instantaneously to solar forcing and communicates this warming through the troposphere by convection. We conclude that the effect of methane thermodynamics on the large-scale circulation plays a primary role in determining the location and timing of shifts of Titan's methane clouds.

There is current observational evidence that the polar clouds have already ceased (10). Although somewhat before our simulations predict this cessation to happen, the actual occurrence makes our results compelling because cessation of polar clouds is a robust feature of our model. In the coming years, Cassini and ground-based observations of Titan's clouds will reveal the seasonal response of the midlatitude clouds. We predict that these clouds will move noticeably in the next 5–7 years if sufficient methane is available at lower latitudes. Future observations by the Cassini mission might also allow us to constrain the distribution of surface sources of methane, which could still be playing a role in determining the observed cloud features.

Methods

We use an axisymmetric dynamical core that solves the primitive equations on a regular grid in the latitude-height plane. In our current model we ignore horizontal mixing of heat, moisture, and momentum by eddies. We use a simplified suite of atmospheric physics including the following: (i) gray infrared radiation by a uniformly mixed absorber, neglecting methane radiative feedback, (ii) a simplified Betts-Miller moist convection scheme that relaxes to specified relative humidity when unstable for our moist simulations, (iii) a bulk aerodynamic surface flux scheme that relaxes methane in the first atmospheric layer toward a specified relative humidity that is calculated in comparison to saturation at the surface temperature and assuming a fixed wind speed, (iv) vertical mixing in the boundary layer with a fixed diffusion coefficient and fixed layer depth (1,500–1,250 millibars), and (v) a simple slab surface with specified heat capacity. We choose to ignore the radiative feedback of methane to study the effect of methane thermodynamics in a clean manner. We also choose to ignore any shortwave absorption in the interior, which primarily occurs in the stratospheric tholin haze, because we are concerned with the troposphere that contains most of the mass of the atmosphere. Gray radiation allows for computation efficiency, and gray radiative equilibrium for a well mixed absorber was found to reproduce the essential features of the observed temperature profile in the troposphere. Although radiative equilibrium can account for the observed temperature lapse rate at Titan's low latitudes (between -1.0 and 0.0 K/km in the lowest

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